Astro <sup>507</sup>Lecture <sup>13</sup>Feb. 24, <sup>2020</sup>

Announcements:

- Preflight 3 due Friday: the CMB!
- Prodigal Instructor returns, thanks for your patience

In the distant past:

 $\overline{\phantom{0}}$ 

Robertson-Walker and relativistic cosmology

- re-derived redshift  $z a$  relation, and cosmic time dilation
- PS2: explored RW metric, introduced "conformal time"

Today: last day of cosmological boot campNext time: apply tools to Dark Energy

#### Recap: Photon Propagation in FLRW

for <sup>a</sup> radial photon (i.e., coming to us)

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$$
d\ell_{\text{com}} = \frac{dr}{\sqrt{1 - \kappa r^2/R^2}} = \frac{dt}{a(t)} = d\eta
$$



Why is  $\eta$  a "conformal" time? conformal transformation <sup>=</sup> angle-preserving  $ds^2=a(\eta)^2$  (d $\eta^2$  $d\ell_{\mathsf{com}}^2) = a(\eta)^2 \times (\mathsf{Minkowski\ form})$  preserves Minkowski "angles" in spacetime $\rightarrow$  lightcones keep straight slopes:  $d\eta/d\ell_{\text{com}} = 1$  on cone

compare photon trajectory in  $(t, \ell_{\mathsf{com}})$  plane: at early times: light cone "slope"  $dt/d\ell_{\mathsf{com}}=a(t)\ll1$  Q: what does this look like? why inconvenient?www: light cones:  $(t, \ell_{\mathsf{com}})$  vs  $(\eta, \ell_{\mathsf{com}})$ plane

# Cosmic Causality



Now RW metric:  $ds^2 = dt^2 - a^2 d\ell_{\rm com}^2$ introduce new time variable  $\eta$ : conformal time defined by  $d\eta = dt/a(t)$  (see PS2)

$$
ds^2 = a(\eta)^2 \left( d\eta^2 - d\ell_{\text{com}}^2 \right)
$$

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Q: implications?

$$
ds^{2} = a(\eta)^{2} \left( d\eta^{2} - d\ell_{\text{com}}^{2} \right) = a(\eta)^{2} \times \text{(Minkowski structure)}
$$

has same features as Minkowski space⇒ *light cones still defined*<br>when use comoving lenc

when use comoving lengths and conformal time



For a flat universe  $(\kappa = 0)$ , it's even better:

$$
ds^{2} = a(\eta)^{2} (d\eta^{2} - dr_{\text{com}}^{2}) = a(\eta)^{2} \times \text{ (exact Minkowski form)}
$$

In either case  $\rightarrow$  spacelike, timelike, lightlike divisions same<br>and in (m  $\ell$  ) space: and in  $(\eta, \ell_{\text{com}})$  space:

light cone structure the same  $\Rightarrow$  causal structure the same!

#### Namely:

- <sup>a</sup> spacetime point can only be influencedby events in past light cone
- <sup>a</sup> spacetime point can only influenceevents in future light cone

So far: like Minkowski

<sub>o</sub> New cosmic twist: finite cosmic age Q: implications for causality?

### Causality: Particle Horizon

past light cone at  $t$  defined by photon propagation over cosmic history:

$$
\int_{t_{\text{em}}=0}^{t_{\text{obs}}=t_{0}} \frac{d\tau}{a(\tau)} = \int_{0}^{r_{\text{em}}} \frac{dr}{\sqrt{1 - \kappa r^{2}/R^{2}}} \equiv d_{\text{hor},\text{com}}(t_{0})
$$

where  $d_{\mathsf{hor},\mathsf{com}}$  is the comoving distance<br>photon has traveled since big bang photon has traveled since big bang

if  $d_{\mathsf{hor},\mathsf{com}} = \int_0^t$  $\mathbf{v}$  $\frac{\partial^{\tau} d\tau}{\partial(\tau)}$  converges then only a finite part of U has affected us  $\rightarrow$   $d_{\sf hor}$  defines *causal boundary* → comoving" particle horizon"

Q: physical implications of <sup>a</sup> particle horizon?

Q: role of finite age?

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Q: sanity check–simple limiting case with obvious result?

### Particle Horizons: Implications

our view of the Universe:  $\star$  astronomical info comes from events along *past light cone*  $\star$  geological info comes from

past world line



- if particle horizon finite (i.e.,  $\neq \infty$ ), then  $d_{\mathsf{horiz,com}}$ :
- **•** gives comoving size of **observable universe**
- encloses region which can communicate over cosmic time→ causally connected region<br>sets "zone of influence" eve
- sets "zone of influence" over which particles can"notice" and/or affect each each other
- and local physical processes can "organize" themselves e.g., shouldn't see bound structures large than particle horizon!  $\overline{\phantom{0}}$

So *is*  $d_{\mathsf{hor}}$  *finite?* depends on details of  $a(t)$  evolution as  $t\rightarrow 0$ : behavior near singularity crucial

will see in PS3: for matter, radiation domination:

- $\bullet$   $d_{\sf hor}$  finite
- and  $d_{\text{hor}} \rightarrow 0$  for  $t \rightarrow 0$   $n_1$  0



Q: implications for CMB?Hint: observed  $T_\mathsf{CMB}(\theta,\phi)$  isotropic to 5th decimal place...

will see in coming weeks

⊲ inflation (if real!) adds twist! <sup>8</sup>

### Cosmic Distance Measures

More examples of how spacetime properties impose relationships among observables

Warmup: Newtonian cosmologyanother sanity check, limiting caseQ: validity range?

Consider Newtonian cosmo:

- $\bullet$  given observed  $z$ , what is distance  $d_{\sf Newt}$ ?
- Q: good for which  $z$ ?
- Q: complications in full FLRW universe?

 $\circ$ 

### "Newtonian Distance"

Newtonian cosmology:

• small speeds, weak gravityignore curvature

Hubble's Law:

$$
H_0 d_{\text{Newt}} \equiv v \simeq cz \tag{1}
$$

applicability:  $z\ll1$ solve:

$$
d_{\text{Newt}} = \frac{c}{H_0} z = d_H z
$$

- naïve distance  $d_{\mathsf{Newt}}$  is linear in  $z$
- it is proportional to the Hubble length  $d_H$ 
	- fraction  $d_{\rm Newt}/d_{\rm H}=z$ ; compare  $t_{\rm lookback}/t_{\rm H}\approx z$

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## Distances and Relativity

Basic but crucial distinction, important to remember:

In *Newtonian/pre-Relativity* physics: space is *absolute* 

- "distance" has unique, well-defined meaning: ⇒ Euclidean separation between points<br>san think of as "intrinsis" to objects ar
- can think of as "intrinsic" to objects and points

In *Special and General Relativity*: space *not* absolute

- distance observer-dependent, not intrinsic to objects, events
- different well-defined measurements can lead todifferent results for distance

In FLRW universe, "distance" not unique: answer depends on

• what you measure

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• how you measure it

### Proper Distance

So far: have constructed *comoving* coordinates which expand with Universe ("home" of fundamental observers)

RW metric: encodes **proper distance** 

i.e., *physical* separations as measured by metersticks/calipers:

⊳ in RW frame i.e., by comoving observers=FOs

 $\triangleright$  at one fixed cosmic instant  $t$ 

$$
d\ell_{\text{prop}}^2 = a(t)^2 d\ell_{\text{com}}^2 = a(t)^2 \left( \frac{dr^2}{1 - \kappa r^2/R^2} + r^2 d\theta^2 + r^2 \sin^2 \theta d\phi^2 \right)
$$

Can read off proper distances for small displacements as measured by FOs at time  $t$ :

- • $\bullet$  dl<sup>prop</sup> =  $= a(t) d\ell_r^{\text{com}} = a(t) dr/\sqrt{1 - \kappa r^2/R^2}$
- • $\begin{array}{l}\n\bullet \ d\ell_{\theta}^{\text{prop}} = a(t) \, d\ell_{\theta}^{\text{com}} = a(t) \, r d\theta \\
\bullet \ d\ell_{\theta}^{\text{prop}} = a(t) \, d\ell_{\theta}^{\text{com}} = a(t) \, r d\theta\n\end{array}$

 $\overline{2}$ 

• 
$$
d\ell_{\phi}^{\text{prop}} = a(t) d\ell_{\phi}^{\text{com}} = a(t) r \sin \theta d\phi
$$
  
Or how to find distance for finite.

Q: how to find distance for finite displacements?

for finite displacements: integrate small ones

e.g., radial distance (at  $t$ ) between  $r=0$  and  $r$  is

$$
\ell_r^{\text{prop}} = a(t)\ell_r^{\text{com}} = a(t) \int_0^r d\zeta / \sqrt{1 - \kappa \zeta^2 / R^2} \tag{2}
$$

Note:  $d\ell_r^{\sf prop}/dt = \dot{a}\,\ell_r^{\sf com} =$  $\rightarrow$  i.e., at a fixed cosmic time t  $= H \ell_r^{\rm prop}$  $r^{p_1o_1p}$  exactly! proper distance increase exactly obeys Hubble Law! Q: what does this mean for points with  $\ell_{\mathsf{r}}^{\mathsf{prop}} > d_H$ ? Q: is this <sup>a</sup> problem?

Q: how would you in practice measure  $\ell_r^{\sf prop}$  for large  $r$ ?

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### Apparent Brightness of <sup>a</sup> Standard Candle

# consider a "**standard candle**"

• object of known rest-frame luminosity

$$
L_{\text{em}} = \frac{dE_{\text{em}}}{dt_{\text{em}}}
$$

- emitting isotropically
- at epoch with  $a_{\text{em}}$  and at rest in cosmic frame
- also, assume no absorbing medium anywhere on sightline

if unresolved = **point source**, observables:<br>1 redabift

1. redshift  $z_{\sf em}$ 

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2. observed flux (apparent brightness)

 $F_{\text{obs}} = dE_{\text{obs}}/dt_{\text{obs}} dA$ 

 $F_{\rm obs}$ *rL*em

summed over all wavelengths: "bolometric"

 $Q$ : Newtonian relation between  $L$  and  $F$ ?



<u>Goal:</u> given std candle  $L_{\mathsf{em}}$ , want to relate observed  $z_{\textsf{em}}$  and  $F_{\textsf{obs}}$ 

 $\Rightarrow$  find expression for luminosity distance<br>closined by Newtonian (Euclidean formula defined by Newtonian/Euclidean formula:

$$
d_{\mathsf{L}}(z_{\mathsf{em}}) \equiv \sqrt{\frac{F_{\mathsf{obs}}}{4\pi L_{\mathsf{em}}}}
$$
(3)

Q: effects in cosmological setting?15

Strategy: start with observation, work back

#### Observation:

FO with telescope, area  $A_{\sf det}$ in time interval  $\delta t_{\rm obs}$ measures total energy  $\delta\mathcal{E}_{\mathsf{obs}}$ , avg photon energy  $\epsilon_{\mathsf{obs}}$ ;

observed flux (bolometric,  $\lambda$ -integrated) given by

$$
\delta \mathcal{E}_{\text{obs}} = F_{\text{obs}} A_{\text{det}} \delta t_{\text{obs}} \tag{4}
$$

 $F_{\mathsf{obs}}$  is rate of energy flow per unit area as measured in observer frame

 $Q:$  what's invariant/observer independent as signal propagates?

Standard candle emitter: luminosity  $L_{\mathsf{em}}$  at  $a_{\mathsf{em}}, z_{\mathsf{em}}$ with average photon energy  $\epsilon_{\mathsf{em}}$ 

- choose  $r_{em} = 0$  as center
- light "cone" (sphere) today reaches us, has present area  $A_{\mathsf{sph}} = 4\pi a_{\mathsf{obs}}^2 r^2 = 4\pi r^2$



key physical principle:

photon counts are invariant

i.e., all observers agree on how many detector registers Q: how to quantify photon number conservation?