Astro 507 Lecture 13 Feb. 24, 2020

Announcements:

- Preflight 3 due Friday: the CMB!
- Prodigal Instructor returns, thanks for your patience

In the distant past:

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Robertson-Walker and relativistic cosmology

- \bullet re-derived redshift z-a relation, and cosmic time dilation
- PS2: explored RW metric, introduced "conformal time"

Today: last day of cosmological boot camp Next time: apply tools to Dark Energy

Recap: Photon Propagation in FLRW

for a radial photon (i.e., coming to us)

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$$d\ell_{\rm com} = \frac{dr}{\sqrt{1 - \kappa r^2/R^2}} = \frac{dt}{a(t)} = d\eta$$

conformal time η_0 η_0 η_0

Why is η a "conformal" time? conformal transformation = angle-preserving $ds^2 = a(\eta)^2 (d\eta^2 - d\ell_{com}^2) = a(\eta)^2 \times (Minkowski form)$ preserves Minkowski "angles" in spacetime \rightarrow lightcones keep straight slopes: $d\eta/d\ell_{com} = 1$ on cone

compare photon trajectory in (t, ℓ_{com}) plane: at early times: light cone "slope" $dt/d\ell_{com} = a(t) \ll 1$ Q: what does this look like? why inconvenient? www: light cones: (t, ℓ_{com}) vs (η, ℓ_{com}) plane

Cosmic Causality



Now RW metric: $ds^2 = dt^2 - a^2 d\ell_{com}^2$ introduce new time variable η : **conformal time** defined by $d\eta = dt/a(t)$ (see PS2)

$$ds^{2} = a(\eta)^{2} \left(d\eta^{2} - d\ell_{\rm com}^{2} \right)$$

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Q: implications?

$$ds^2 = a(\eta)^2 \left(d\eta^2 - d\ell_{\text{com}}^2 \right) = a(\eta)^2 \times \text{ (Minkowski structure)}$$

has same features as Minkowski space \Rightarrow *light cones still defined*

when use comoving lengths and conformal time



For a flat universe ($\kappa = 0$), it's even better:

$$ds^2 = a(\eta)^2 \left(d\eta^2 - dr_{\rm com}^2 \right) = a(\eta)^2 \times \text{(exact Minkowski form)}$$

In either case \rightarrow spacelike, timelike, lightlike divisions same and in (η, ℓ_{com}) space:

light cone structure the same \Rightarrow *causal structure the same*!

Namely:

- a spacetime point can only be influenced by events in past light cone
- a spacetime point can only influence events in future light cone

So far: like Minkowski

 ¬ New cosmic twist: finite cosmic age
 Q: implications for causality?

Causality: Particle Horizon

past light cone at t defined by photon propagation over cosmic history:

$$\int_{t_{em}=0}^{t_{obs}=t_0} \frac{d\tau}{a(\tau)} = \int_0^{r_{em}} \frac{dr}{\sqrt{1 - \kappa r^2/R^2}} \equiv d_{hor,com}(t_0)$$

where $d_{\rm hor,com}$ is the comoving distance photon has traveled since big bang

if $d_{\text{hor,com}} = \int_0^t d\tau / a(\tau)$ converges then only a finite part of U has affected us $\rightarrow d_{\text{hor}}$ defines *causal boundary* \rightarrow comoving "particle horizon"

Q: physical implications of a particle horizon?

- Q: role of finite age?
- *Q:* sanity check–simple limiting case with obvious result?

Particle Horizons: Implications

our view of the Universe:
* astronomical info comes from events along past light cone
* geological info comes from

past world line



- if particle horizon finite (i.e., $\neq \infty$), then $d_{\text{horiz,com}}$:
- gives comoving size of observable universe
- encloses region which can communicate over cosmic time \rightarrow causally connected region
- sets "zone of influence" over which particles can "notice" and/or affect each each other
- and local physical processes can "organize" themselves e.g., shouldn't see bound structures large than particle horizon!

So is d_{hor} finite? depends on details of a(t) evolution as $t \rightarrow 0$: behavior near singularity crucial



Hint: observed $T_{CMB}(\theta, \phi)$ isotropic to 5th decimal place...

will see in coming weeks

∞ ⊳ inflation (if real!) adds twist!

Cosmic Distance Measures

More examples of how spacetime properties impose relationships among observables

Warmup: Newtonian cosmology another sanity check, limiting case *Q: validity range?*

Consider Newtonian cosmo:

- given observed z, what is distance d_{Newt} ?
- Q: good for which z?
- *Q*: complications in full FLRW universe?

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"Newtonian Distance"

Newtonian cosmology:

 small speeds, weak gravity ignore curvature

Hubble's Law:

$$H_0 d_{\mathsf{Newt}} \equiv v \simeq cz \tag{1}$$

applicability: $z \ll 1$ solve:

$$d_{\mathsf{Newt}} = \frac{c}{H_0} z = d_H z$$

- naïve distance d_{Newt} is *linear in z*
- $_{\rm H}$ $\,$ it is proportional to the Hubble length $d_{\rm H}$
 - fraction $d_{\text{Newt}}/d_{\text{H}} = z$; compare $t_{\text{lookback}}/t_{\text{H}} \approx z$

Distances and Relativity

Basic but crucial distinction, important to remember:

In *Newtonian/pre-Relativity* physics: space is *absolute*

- "distance" has unique, well-defined meaning:
 ⇒ Euclidean separation between points
- can think of as "intrinsic" to objects and points

In Special and General Relativity: space not absolute

- distance observer-dependent, not intrinsic to objects, events
- different well-defined measurements can lead to different results for distance

In FLRW universe, "distance" not unique: answer depends on

• what you measure

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• how you measure it

Proper Distance

So far: have constructed *comoving* coordinates which expand with Universe ("home" of fundamental observers)

RW metric: encodes proper distance

i.e., *physical* separations as measured by metersticks/calipers:
in RW frame i.e., by comoving observers=FOs

▷ at one fixed cosmic instant t

$$d\ell_{\rm prop}^2 = a(t)^2 d\ell_{\rm com}^2 = a(t)^2 \left(\frac{dr^2}{1 - \kappa r^2/R^2} + r^2 d\theta^2 + r^2 \sin^2\theta d\phi^2\right)$$

Can read off proper distances for small displacements as measured by FOs at time t:

- $d\ell_r^{\text{prop}} = a(t) d\ell_r^{\text{com}} = a(t) dr / \sqrt{1 \kappa r^2 / R^2}$
- $d\ell_{\theta}^{\mathsf{prop}} = a(t) d\ell_{\theta}^{\mathsf{com}} = a(t) r d\theta$

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$$d\ell_{\phi}^{\text{prop}} = a(t) d\ell_{\phi}^{\text{com}} = a(t) r \sin \theta d\phi$$

Q: how to find distance for finite displacements?

for finite displacements: integrate small ones

e.g., radial distance (at t) between r = 0 and r is

$$\ell_r^{\text{prop}} = a(t)\ell_r^{\text{com}} = a(t)\int_0^r d\zeta/\sqrt{1-\kappa\zeta^2/R^2}$$
(2)

Note: $d\ell_r^{\text{prop}}/dt = \dot{a} \ell_r^{\text{com}} = H \ell_r^{\text{prop}}$ exactly! \rightarrow i.e., at a fixed cosmic time t proper distance increase exactly obeys Hubble Law! Q: what does this mean for points with $\ell_r^{\text{prop}} > d_H$? Q: is this a problem?

Q: how would you in practice measure ℓ_r^{prop} for large r?

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Apparent Brightness of a Standard Candle

consider a "standard candle"

• object of known rest-frame luminosity

$$L_{\rm em} = \frac{dE_{\rm em}}{dt_{\rm em}}$$

- emitting isotropically
- at epoch with a_{em} and at rest in cosmic frame
- also, assume no absorbing medium anywhere on sightline

if unresolved = **point source**, observables:

1. redshift zem

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2. observed flux (apparent brightness)

 $F_{\rm obs} = dE_{\rm obs}/dt_{\rm obs} \ dA$

 $F_{\rm obs}$

- summed over all wavelengths: "bolometric"
- Q: Newtonian relation between L and F?



Goal: given std candle $L_{\rm em}$, want to relate observed $z_{\rm em}$ and $F_{\rm obs}$

⇒ find expression for luminosity distance defined by Newtonian/Euclidean formula:

$$d_{\rm L}(z_{\rm em})\equiv \sqrt{rac{F_{\rm obs}}{4\pi L_{\rm em}}}$$

(3)

G = Q: effects in cosmological setting?

Strategy: start with observation, work back

Observation:

FO with telescope, area A_{det} in time interval δt_{obs} measures total energy $\delta \mathcal{E}_{obs}$; avg photon energy ϵ_{obs}

observed flux (bolometric, λ -integrated) given by

$$\delta \mathcal{E}_{\rm obs} = F_{\rm obs} A_{\rm det} \delta t_{\rm obs} \tag{4}$$

 $F_{\rm obs}$ is rate of energy flow per unit area as measured in observer frame

Q: what's invariant/observer independent as signal propagates?

Standard candle emitter: luminosity L_{em} at a_{em}, z_{em} with average photon energy ϵ_{em}

- choose $r_{\rm em} = 0$ as center
- light "cone" (sphere) today reaches us, has present area $A_{\rm sph}=4\pi a_{\rm obs}^2 r^2=4\pi r^2$



key physical principle:

photon counts are invariant

i.e., all observers agree on how many detector registers *Q: how to quantify photon number conservation?*